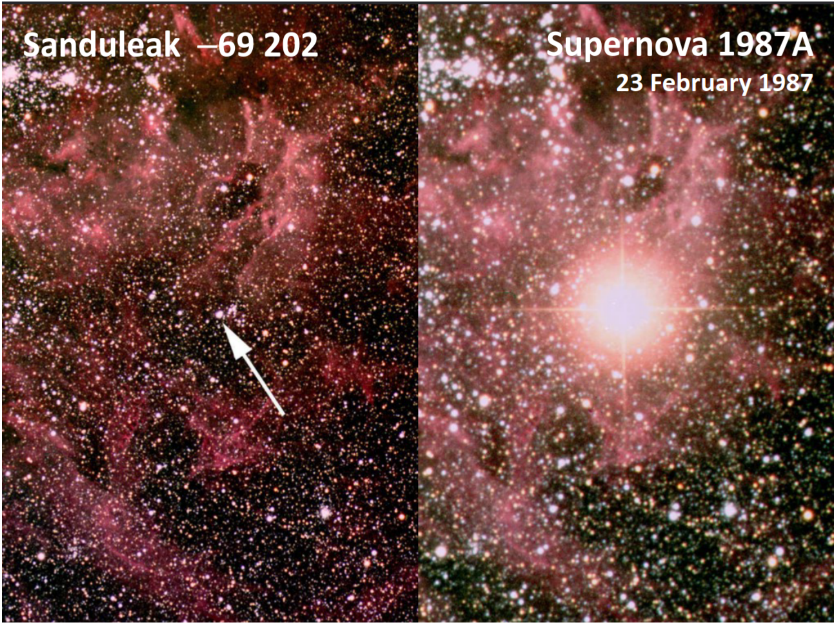
LNV Neutrino Self-Interaction in Core-Collapse Supernovae

Anna M. Suliga
with Fuller, Graf, Cheong, Froustey, Shalgar, Kherer, Scholer
arXiv: 2410.01080

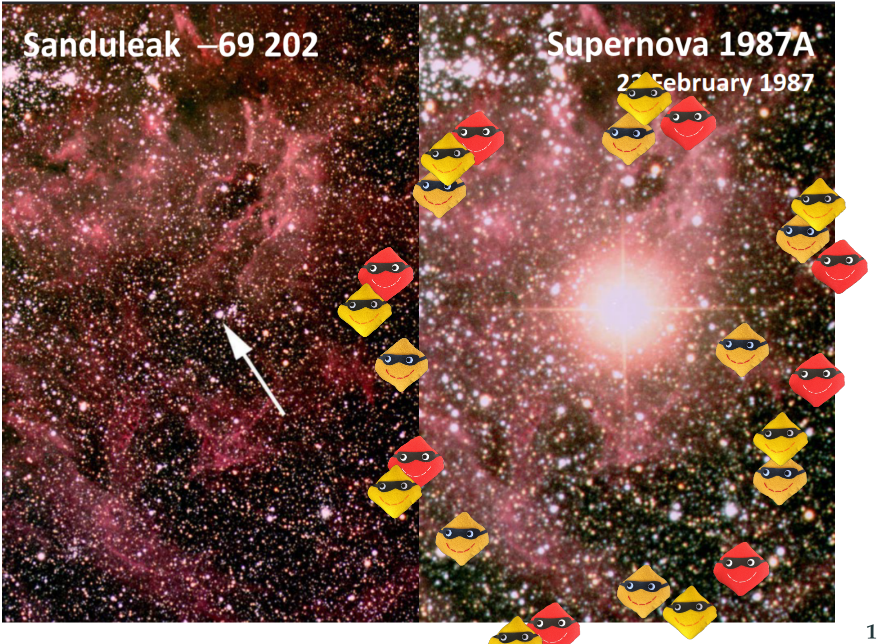
Neutrino Theory Network Fellow
New York University

Mitchell Institute,
Texas A&M University
May 15, 2025

Why is studying astrophysical neutrinos crucial?



Why is studying astrophysical neutrinos crucial?



Why is studying astrophysical neutrinos crucial?

Hubble (2017)

JWST (2023)

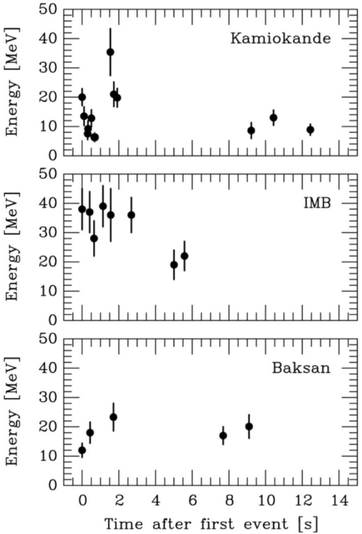
- Neutron star remnant

  Fransson et al. (2024)

- Binary system

  Morris & Podsiadlowski (2007), (2009)

Established track record of neutrino discoveries: SN 1987A



Courtesy of G. Raffelt

- Neutrino detection from SN 1987A:
  - confirmed the core-collapse scenario
  - 99% of the energy emitted in neutrinos
  - best limit at the time on the ν mass

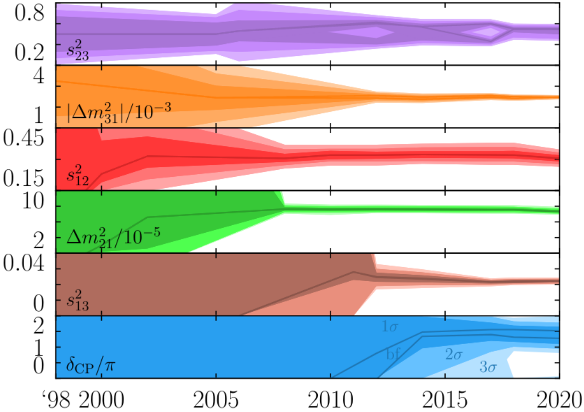
Towards Precise Neutrino Properties Measurements

We known now:

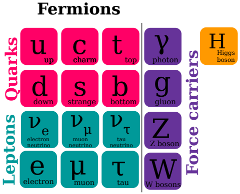
- large mixing angles
- non-zero masses

Remaining questions

- Majorana vs Dirac
- absolute masses
- degree of CP violation



Snowmass Neutrino Frontier Report (2022)



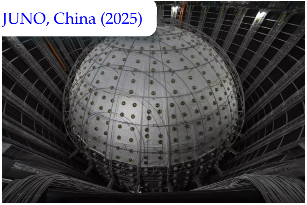
$$\begin{bmatrix}\nu_e\\ \nu_\mu\\ \nu_\tau\end{bmatrix} = U \begin{bmatrix}\nu_1\\ \nu_2\\ \nu_3\end{bmatrix}$$

beam, atmospheric — **beam, reactor** — **solar, reactor**

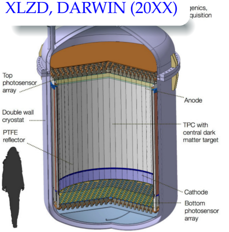
$$U = \begin{bmatrix}1 & 0 & 0\\ 0 & c_{23} & s_{23}\\ 0 & -s_{23} & c_{23}\end{bmatrix}\begin{bmatrix}c_{13} & 0 & s_{13}e^{-i\delta_{CP}}\\ 0 & 1 & 0\\ -s_{13}e^{i\delta_{CP}} & 0 & c_{13}\end{bmatrix}\begin{bmatrix}c_{12} & s_{12} & 0\\ -s_{12} & c_{12} & 0\\ 0 & 0 & 1\end{bmatrix}$$

How to achieve full picture of neutrinos? All hands on deck!

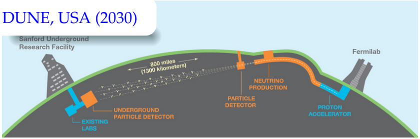
JUNO, China (2025)



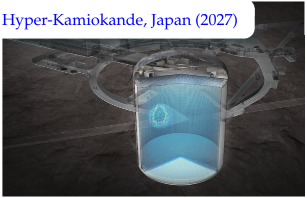
XLZD, DARWIN (20XX)



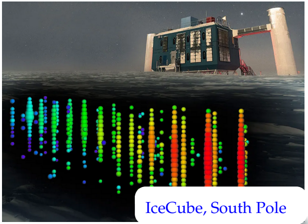
DUNE, USA (2030)



Hyper-Kamiokande, Japan (2027)



Rubin Observatory, Chile (2025)



IceCube, South Pole

- Many new experiments comimg online soon
- variety of approaches $\rightarrow$ superb sensitivity

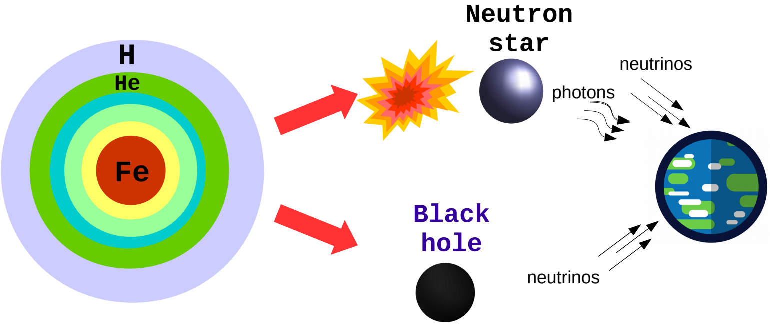
- Complementarity with:
  - reactor and accelerator searches
  - electromagnetic surveys
  - other astrophysical messengers

Neutrinos from Core-collapse Supernovae

Why are neutrinos important for a core-collapse supernova?

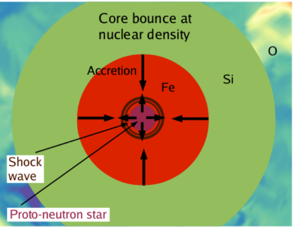
Neutrinos:

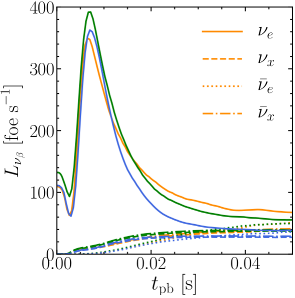
- $\sim 10^{58}$ of them emitted from a single core collapse
- only they can reveal the deep interior conditions
- only particles detectable from the collapse to a black hole



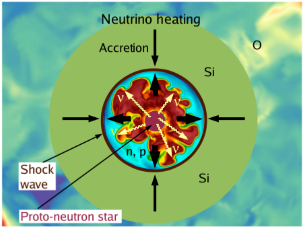
Different Phases of Supernova Explosion

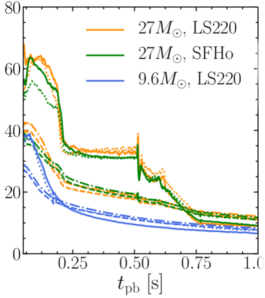
- Infall phase, ν_e burst ~ 40 ms



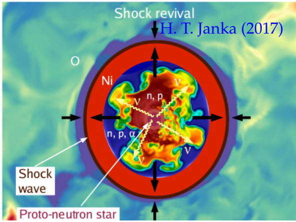


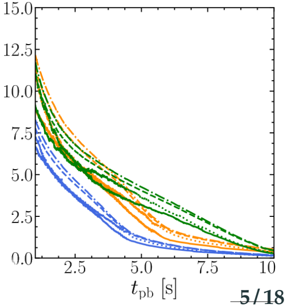
- Accretion phase, ~ 100 ms





- Cooling phase, ~ 10 s





Why core-collapse supernovae are good physics probes?

Advantages

- extreme physical conditions not accessible on Earth
- within the reach of existing and upcoming detectors

What can we learn with a variety of detectors?

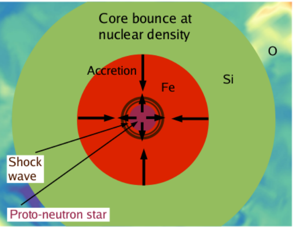
- explosion mechanism — Bethe & Wilson (1985), Fischer et al. (2011)...
- nucleosynthesis — Woosley et al. (1994), Surman & McLaughlin (2003)...
- compact object formation — Warren et al. (2019), Li, Beacom et al. (2020)...
- neutrino mixing — Balantekin & Fuller (2013), Tamborra & Shalgar (2020)...
- non-standard physics — McLaughlin et al. (1999), de Gouvêa et al. (2019) ...

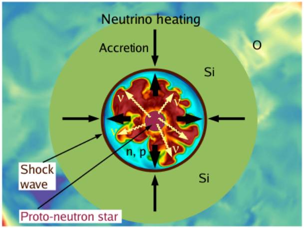
Neutrinos from Supernovae as Probes of New Physics

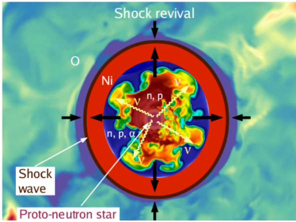
Different Phases of Supernova Explosion

- Infall phase, ν_e burst ~ 40 ms
- Accretion phase, ~ 100 ms
- Cooling phase, ~ 10 s

H. T. Janka (2017)







New neutrino physics affects the core-collapse supernovae:

- change diffusion time → possible change in the star's fate
- changed diffusion time → changed duration of the neutrino signal
- new cooling channel → affects explosion probability

astrophysical feedback often ignored

Which bounds remain unchanged with astrophysical feedback?

Do Neutrinos Have Self-Interactions?

IL NUOVO CIMENTO VOL. XXXIII, N. 5 1° Settembre 1964

Do Neutrinos Interact between Themselves?

Z. BIALYNICKA-BIRULA

Institute of Physics, Polish Academy of Sciences - Warsaw

(ricevuto il 26 Giugno 1964)

1. – Introduction.

The neutrino is the only elementary particle, which, according to our present knowledge, does not take part in other than weak and gravitational interactions. Its role in nature is not yet fully understood and its interaction properties are only partially known.

The purpose of this note is to answer the following question: Do the present experimental data allow for the existence of interactions between neutrinos much stronger than their weak interactions? The answer to this question is positive. It turns out that such interactions even if they were 10^6 times stronger than weak interactions could not be detected with the present experimental accuracy.

Zofia Białynicka-Birula (1964)

Lepton number violating neutrino self-interactions

Motivation - to be taken with a grain of salt:

- lepton number conservation - accidental symetry
- potential cosmological hints

  Barenboim et al. (2019), Song, Gonzalez-Garcia, Salvado (2018), ..
- strong impact on core-collapse supernova

  Kolb et al. (1982), Fuller et al. (1988), Farzan et al. (2018), **AMS**, Tamborra (2020), ...

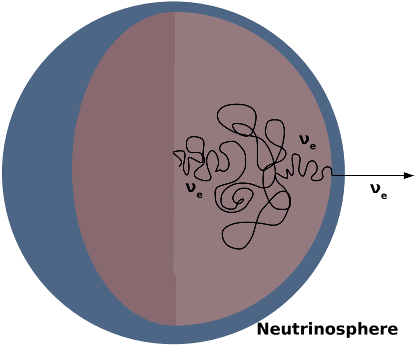
New Interaction Lagrangian

$$\mathcal{L}^{\phi} = g_{\phi,\alpha\beta}\, \phi\, \overline{\nu_{L,\alpha}}\, \nu^{c}_{L,\beta}$$

Probability of the New Interaction

$$\sigma_{\nu\mathrm{SI}} \approx \frac{G^2_{\nu\mathrm{SI}}}{8\pi} E^1_\nu E^2_\nu (1 - \cos\theta)$$

Neutrino Trapping and $\beta-$equilibrium



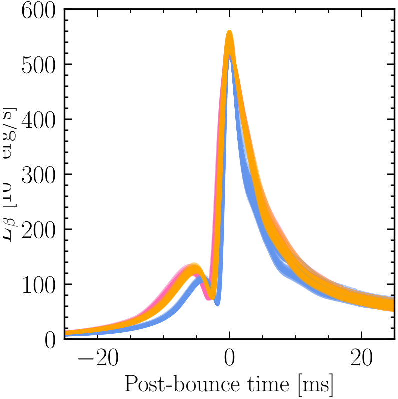
Neutrino trapping

$$A(N,Z) + \nu \rightarrow A(N,Z) + \nu$$

β-equilibrium

$$e^- + p \rightleftharpoons \nu_e + n$$

$$e^+ + n \rightleftharpoons \bar{\nu}_e + p$$



Neutrino trapping

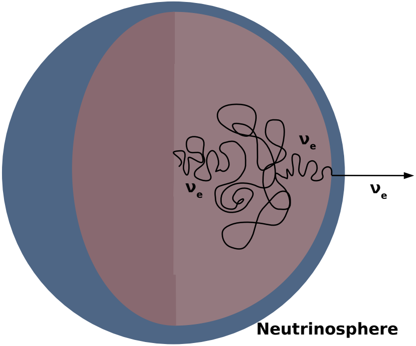
$$A(N,Z) + \nu \to A(N,Z) + \nu$$

β-equilibrium

$$e^- + p \rightleftharpoons \nu_e + n$$

$$e^+ + n \rightleftharpoons \bar{\nu}_e + p$$

Neutrino Trapping and $\beta-$equilibrium



Neutrino trapping

$$A(N,Z)+\nu \to A(N,Z)+\nu$$

β-equilibrium

$$e^- + p \rightleftharpoons \nu_e + n$$

$$e^+ + n \rightleftharpoons \bar{\nu}_e + p$$

LNV νSI Implementation:

Thermalize the population of ν and $\bar{\nu}$ once $\rho \sim 10^{11} - 10^{12}\text{g cm}^{-3}$

$\nu_e \rightleftharpoons \nu_e, \bar{\nu}_e, \nu_\mu, \bar{\nu}_\mu, \nu_\tau, \bar{\nu}_\tau,$ $\quad$ $\nu_e \rightleftharpoons \nu_e, \bar{\nu}_e, \nu_x, \bar{\nu}_x,$ $\quad$ $\nu_e \rightleftharpoons \nu_e, \bar{\nu}_e$

Static, Homogenous and Isotropic Boltzmann Equation

Boltzmann Equation

$$\frac{df_\nu}{dt} = (1 - f_\nu)j_\nu - f_\nu \chi_\nu \ ,$$

Electron fraction evolution - weak rates

$$e^- + p \leftrightharpoons \nu_e + n$$

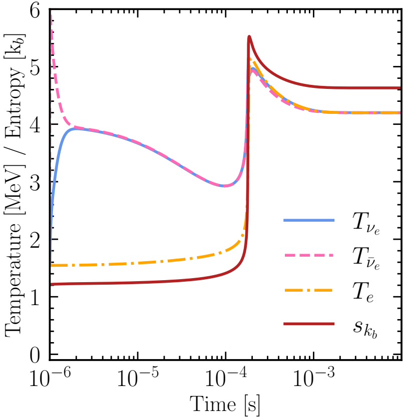
$$e^+ + n \leftrightharpoons \bar{\nu}_e + p$$

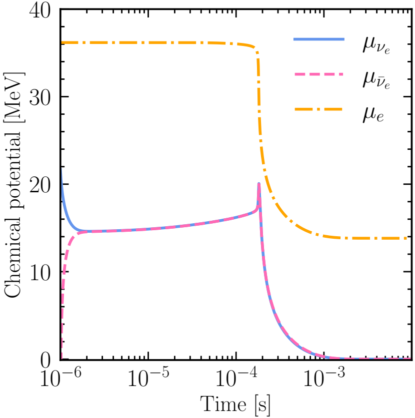
$$\frac{dY_e}{dt} = R_{\nu_e} - R_{\bar{\nu}_e} - R_{e^-} + R_{e^+} \ ,$$

Temperature and chemical potential evolution for leptons

$$\frac{dT_i}{dt} = \left(\frac{\partial \rho_i}{\partial \mu_i} \frac{dn_i}{dt} - \frac{\partial n_i}{\partial \mu_i} \frac{d\rho_i}{dt} \right) \Big/ \left(\frac{\partial n_i}{\partial T_i} \frac{\partial \rho_i}{\partial \mu_i} - \frac{\partial n_i}{\partial \mu_i} \frac{\partial \rho_i}{\partial T_i} \right) \ ,$$

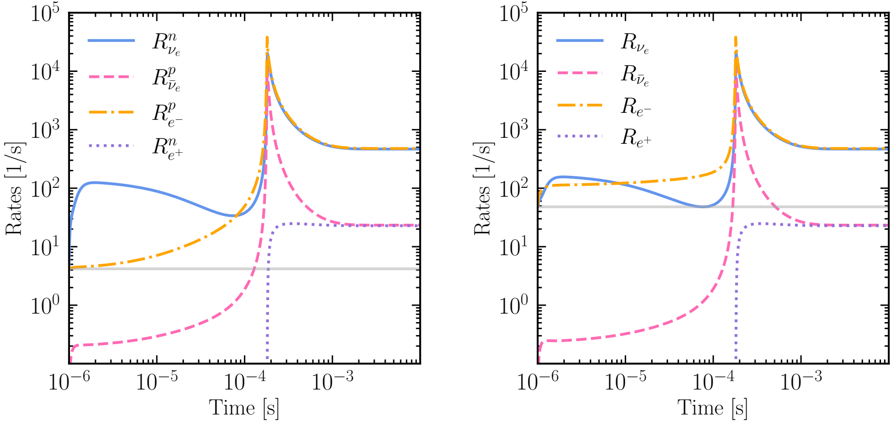
$$\frac{d\mu_i}{dt} = \left(\frac{\partial \rho_i}{\partial T_i} \frac{dn_i}{dt} - \frac{\partial n_i}{\partial T_i} \frac{d\rho_i}{dt} \right) \Big/ \left(\frac{\partial n_i}{\partial \mu_i} \frac{\partial \rho_i}{\partial T_i} - \frac{\partial n_i}{\partial T_i} \frac{\partial \rho_i}{\partial \mu_i} \right) \ .$$





- new interactions quickly equilbrate ν_e and $\bar{\nu}_e$ seas
- enhanced ν_e and e^- captures heat up the matter
- similar results for all flavors equilibration

Weak reaction rates



- initial increase in $\nu_e + n$, $\nu_e + A$ and $e^- + A$
- enhanced ν_e and e^- captures heat up the matter
- similar results for all flavors equilibration

Fast LNV νSI - Approximate Evolution

LNV νSI timescale much faster than weak timescale $\rightarrow$ a single ν species evolution

$$\sum_\alpha \left(\frac{dn_{\nu_\alpha}}{dt} + \frac{dn_{\bar{\nu}_\alpha}}{dt}\right) = \frac{\delta n_\nu}{\delta t}$$

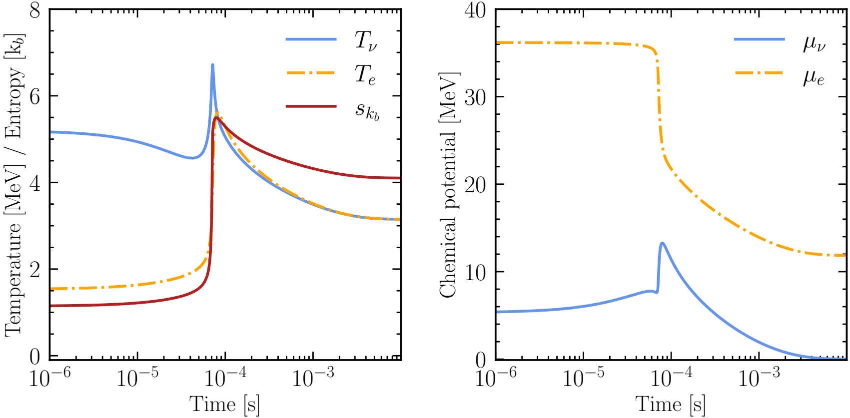
$$\sum_\alpha \left(\frac{d\rho_{\nu_\alpha}}{dt} + \frac{d\rho_{\bar{\nu}_\alpha}}{dt}\right) = \frac{\delta \rho_\nu}{\delta t}$$

sum over charged-current weak interactions

$$\frac{dT_\nu}{dt} = \frac{\frac{\partial\rho_\nu}{\partial\mu_\nu}\frac{\delta n_\nu}{\delta t} - \frac{\partial n_\nu}{\partial\mu_\nu}\frac{\delta\rho_\nu}{\delta t}}{2N_F\left(\frac{\partial n_\nu}{\partial T_\nu}\frac{\partial\rho_\nu}{\partial\mu_\nu} - \frac{\partial n_\nu}{\partial\mu_\nu}\frac{\partial\rho_\nu}{\partial T_\nu}\right)}$$

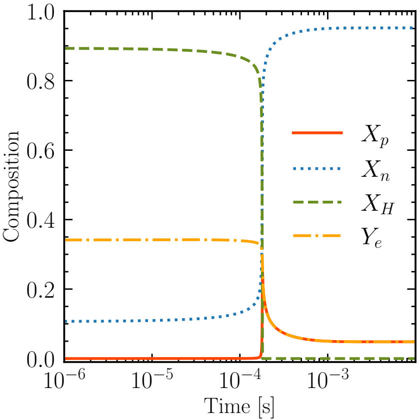
$$\frac{d\mu_\nu}{dt} = \frac{\frac{\partial\rho_\nu}{\partial T_\nu}\frac{\delta n_\nu}{\delta t} - \frac{\partial n_\nu}{\partial T_\nu}\frac{\delta\rho_\nu}{\delta t}}{2N_F\left(\frac{\partial n_\nu}{\partial\mu_\nu}\frac{\partial\rho_\nu}{\partial T_\nu} - \frac{\partial n_\nu}{\partial T_\nu}\frac{\partial\rho_\nu}{\partial\mu_\nu}\right)}$$

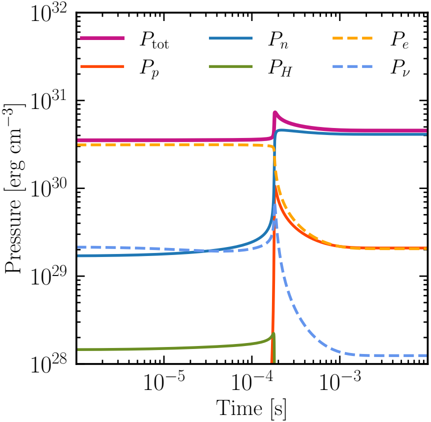
Evolution of Thermodynamical Quantities



- the same qualitative results for all six flavor equilibration

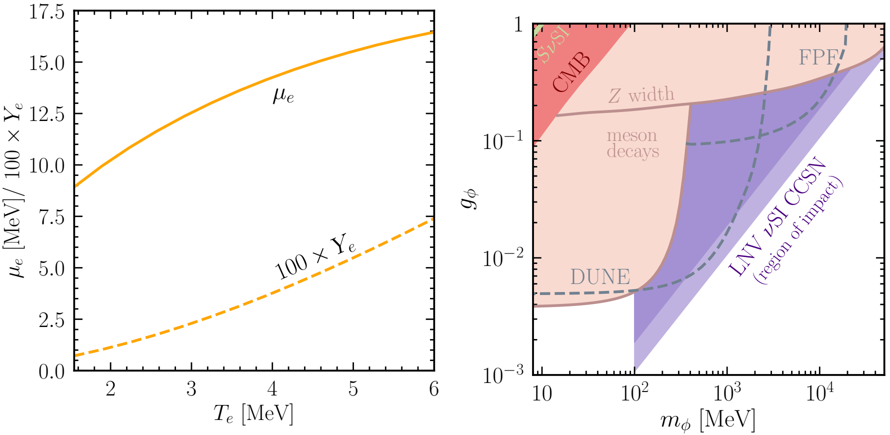
Composition and Pressure Support of the Core





- s_{k_b} - entropy generation shifts composition towards no heavy nuclei
  $X_H \propto s_{k_\mathrm{B}}^{1-\langle A \rangle} n_p^Z n_n^N \exp(E_b/T_e)$
- enhanced deleptonization changes the pressure support of the core

New β-equilibrium with LNV νSI



- regardless of the final T_e the new equilibrium has a very low Y_e
  $\mu_e = \delta m_{np} - T_e \ln\left(\frac{Y_e}{1-Y_e}\right)$, with $Y_e = \frac{1}{\pi^2 \rho} \int_0^\infty dp_e\, p_e^2\, f_e(E_e, T_e, \mu_e)$
- complementarity with future accelerator-based experiments

Conclusions

Strong LNV νSI

- can significantly affect the supernova core evolution
- can cause a multi-flavor neutrino burst
- overlap with the future terrestrial probes

Core-collapse supernovae

- serve as testing grounds in constraining standard and new physics
- reliable limits, only when the sources are accurately modeled

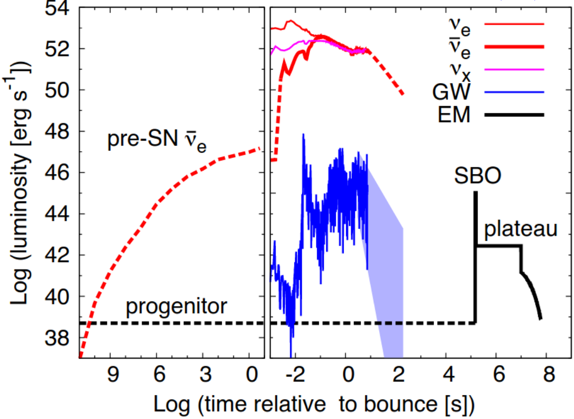
Exciting times ahead

Thank you for the attention!

Backup

Core-Collapse Supernova Light Curve

Nakamura et al. (2016)



Partial Derivatives for the Fermi-Dirac distributions

The partial derivatives for the Fermi-Dirac distributions are given by Escudero (2020)

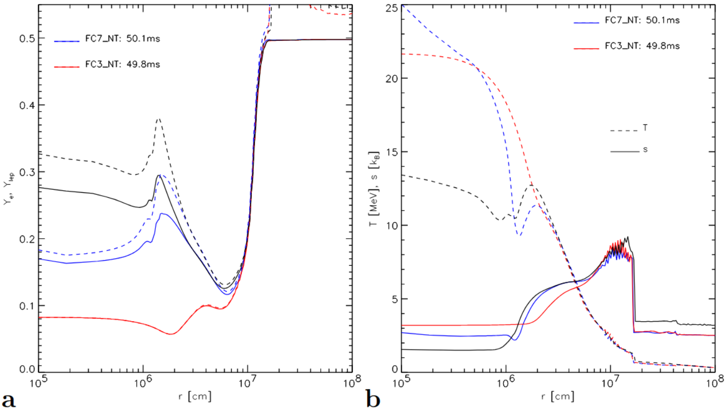
$$\frac{\partial n}{\partial T} = \frac{g}{2\pi^2} \int_m^\infty dE\, E\sqrt{E^2 - m^2}\, \frac{(E-\mu)}{4T^2} \cosh^{-2}\left(\frac{E-\mu}{2T}\right), \quad \text{(1a)}$$

$$\frac{\partial \rho}{\partial T} = \frac{g}{2\pi^2} \int_m^\infty dE\, E^2\sqrt{E^2 - m^2}\, \frac{(E-\mu)}{4T^2} \cosh^{-2}\left(\frac{E-\mu}{2T}\right), \quad \text{(1b)}$$

$$\frac{\partial n}{\partial \mu} = \frac{g}{2\pi^2} \int_m^\infty dE\, E\sqrt{E^2 - m^2}\, \left[2T\cosh\left(\frac{E-\mu}{T}\right) + 2T\right]^{-1}, \quad \text{(1c)}$$

$$\frac{\partial \rho}{\partial \mu} = \frac{g}{2\pi^2} \int_m^\infty dE\, E^2\sqrt{E^2 - m^2}\, \left[2T\cosh\left(\frac{E-\mu}{T}\right) + 2T\right]^{-1} \quad \text{(1d)}$$

Proxy "Internal Deleptonization"



$$\nu_e + N \leftrightarrow \bar{\nu}_e + N$$

M. Rampp et al. (2002)

LS 220 Equation of State: impact of α particles